

# The nuclear cluster of the Milky Way: total mass and luminosity

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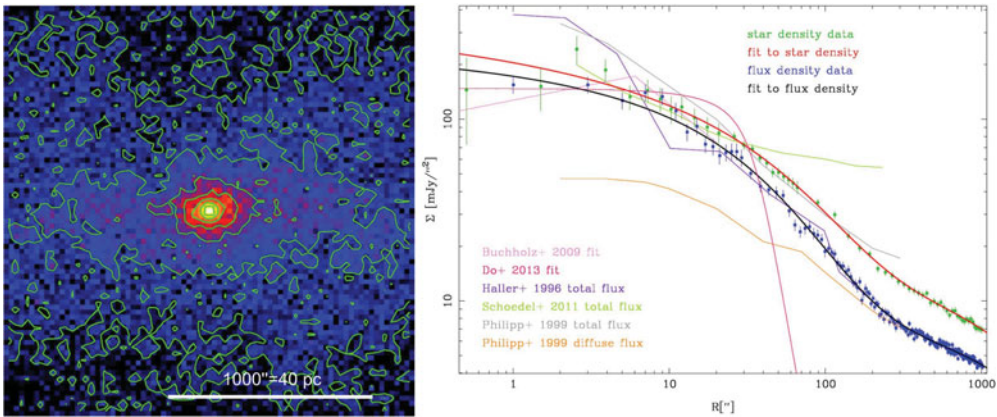
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**Abstract.** Here we present the fundamental properties of the nuclear cluster of the Milky Way. First, we derive its structural properties by constructing a density map of the central 1000'' using extinction-corrected star counts. We can describe the data with a two-component model built from Sersic profiles. The inner nearly spherical component is the nuclear cluster. The outer, strongly flattened component can be identified with the stellar component of the circumnuclear zone. Second, we enlarge the radius inside which detailed dynamics are available from 1 pc to 4 pc. We use more than 10000 individual proper motions and more than 2700 radial velocities. We determine the cluster mass by means of isotropic spherical Jeans modeling. We get a nuclear cluster mass within 100'' of  $M_{100''} = (6.11 \pm 0.52|_{\text{fix } R_0} \pm 0.97|_{R_0}) \times 10^6 M_{\odot}$ , which corresponds to a total cluster mass of  $M_{\text{NC}} = (13.08 \pm 2.51|_{\text{fix } R_0} \pm 2.08|_{R_0}) \times 10^6 M_{\odot}$ . By combination of our mass with the flux we calculate  $M/L = 0.50 \pm 0.12 M_{\odot}/L_{\odot, K_s}$  for the central 100''. That is broadly consistent with a Chabrier IMF. With its mass and a luminosity of  $M_{K_s} = -15.30 \pm 0.26$  the nuclear cluster is a bright and massive specimen with a typical size.

## 1. Introduction

In the centers of many late-type galaxies one finds massive stellar clusters: the nuclear star clusters (Phillips *et al.* 1996, Matthews & Gallagher 1997, Carollo *et al.* 2002). The nuclear clusters are central light overdensities on a scale of about 5 pc (Böker *et al.* 2004). Also, the central light concentration of the Milky Way (Becklin & Neugebauer 1968) is a nuclear star cluster (Philipp *et al.* 1999, Launhardt *et al.* 2002). Nuclear clusters are comparable in density to globular clusters, but are typically more massive (Walcher *et al.* 2005).

Due to the proximity of the Galactic center (GC), the nuclear cluster of the Milky Way can be observed in much higher detail than any other nuclear cluster (Genzel *et al.* 2010). It is useful to shed light on the properties of nuclear clusters and their formation in general. Here and in Fritz *et al.* (submitted) we concentrate on the most basic properties, mass and luminosity. For that aim we obtain the star distribution of the central galaxy out to 1000'', which allows us to cover far more than the nuclear cluster. In addition, we measure for the first time motions in all three dimensions over a large part of the cluster expanding on Lindqvist *et al.* (1992), Genzel *et al.* (1996), Trippe *et al.* (2008) and Schödel *et al.* (2009).



**Figure 1.** Star distribution in the GC. **Left panel:** Star density map around Sgr A\* ( $r=1000''$  box) with smoothed contours. **Right panel:** Profiles from our data in comparison with the literature. All curves besides our flux density have been scaled, since they do not have the same units or are not extinction-corrected. The thick lines show our fits with each consisting of two independent  $\gamma$ -models. [A COLOR VERSION IS AVAILABLE ONLINE.]

## 2. Density profile and luminosity

For the density maps we combine ( $J/H$ )/ $K_s$ -band data from NACO, WFC3/IR and VISTA. The NACO map provides the high resolution data in the center while the VISTA map gives the largest field of view. With this combination, we have sufficient resolution nearly everywhere for azimuthal decomposition. We use the data in two different ways: first, for number counts of bright stars; and second, for the integrated flux. All data are corrected for extinction by using two color information and masking out infrared dark clouds. We also remove the contribution of the young (OB and IRS7) stars in order to exclude the most recent starburst, whose radial profile is different from the older stars (Bartko *et al.* 2010, Blum *et al.* 2003, Pfuhl *et al.* 2011).

The resulting map is presented in Figure 1. It is apparent that the star distribution is close to circular in the center and more elongated further out. This strengthens the finding of Launhardt *et al.* (2002) that the star-forming nucleus of the Milky Way consists of two components, one which is the close to circular nuclear cluster and a second, which is a flattened nuclear disk. The latter is the stellar component of the circumnuclear disk. To differentiate between these components, we fit them with two Sersic profiles (using GALFIT). We obtain for the nuclear cluster an axis ratio of  $1.10 \pm 0.07$ , a half light radius of  $110 \pm 10'' = 4.4$  pc, a Sersic index of  $1.42 \pm 0.03$  and a total luminosity of  $M_{K_s} = -15.3 \pm 0.26$ . The main uncertainties include the modeling of the incompletely-covered outer component as well as the extinction correction.

We present the flux and number counts radial profile in Figure 1. They are not in perfect agreement with one another, yet, there is similarity between the two and they appear to agree roughly with Becklin & Neugebauer (1968), Haller *et al.* (1996) and Philipp *et al.* (1999). We as well as the aforementioned studies seem to be in disagreement with the recent profile of Schödel (2011), who measured a larger background flux level possible due to missing sky subtraction. Our profile also deviates strongly from the indirect, dynamic profile determination of Do *et al.* 2013. In comparison to ours, their profile has a too large core and a too steep of an outside decline. Note that we fit our profiles for the mass modeling with a combination of two spherical  $\gamma$ -models (Dehnen 1993). By using both, flux and number counts, we include density profile uncertainties in our error budgets.

### 3. Mass

Additionally, to obtain the mass, we measure star velocities in all three dimensions. We employ mainly the NACO data to generate more than 10000 proper motions out to  $80''$ . For radial velocities, we use SINFONI to get more than 2500 radial velocities out to  $80''$  and to go beyond this region, we add about 200 maser velocities from Lindqvist *et al.* (1992) and Deguchi *et al.* (2004). To remove the young stars from our velocity-selected sample, we employ spectra.

Our data show no sign of radial anisotropy out to at least  $40''$ , in contrast to the extrapolation of Do *et al.* (2013). Beyond this, the azimuthal coverage is too incomplete to distinguish radial anisotropy from other deviations from isotropy like flattening. The difference between the dispersion in and perpendicular to the galactic plane (Trippe *et al.* 2008) is explained by flattening of the nuclear cluster in that direction (see Chatzopoulos *et al.*, in prep). We confirm the rotation of the cluster via our radial velocity measurements. The rotation velocity is, however, on the small side of McGinn *et al.* (1989) and Trippe *et al.* (2008) outside of 1 pc.

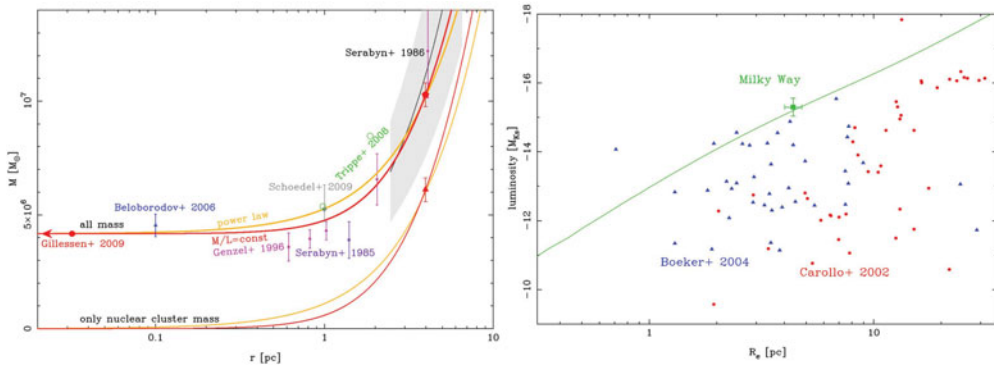
We obtain the mass by Jeans modeling. Here, we restrict ourselves to isotropic spherical symmetric modeling. The mass is modeled by the supermassive black hole (SMBH) in the center and the extended nuclear cluster mass. The SMBH mass is fixed to the black hole distance relation of Gillessen *et al.* (2009). Since the exclusion of flattening in the model a statistical parallax distance would lead to biasing, we use the independent measurement of Gillessen *et al.* 2013 ( $R_0 = 8.2 \pm 0.34$  kpc). For the nuclear cluster we use two models: a power law mass profile and a constant mass to light ratio. We obtain a nuclear cluster mass of  $(6.11 \pm 0.52|_{\text{fix}R_0} \pm 0.97|_{R_0}) \times 10^6 M_\odot$  within  $100''$  which is consistent with most of the literature (Figure 2). The distance-independent mass error is dominated by the profile difference between star counts and light density. The error does not include deviations from our model like flattening. Using the two-dimensional decomposition of Section 2 we obtain a total nuclear cluster mass of  $(13.08 \pm 2.51|_{\text{fix}R_0} \pm 2.08|_{R_0}) \times 10^6 M_\odot$ .

Employing both the mass and luminosity, we find a mass to light ratio of  $M/L = 0.50 \pm 0.12 M_\odot/L_{\odot, K_s}$ . By combining the GC star formation history (Pfuhl *et al.* 2011, Blum *et al.* 2003) with the canonical Kroupa/Chabrier IMF, we obtain  $M/L \approx 0.68$ , as expected value for the cluster. The simplicity of our mass model could possibly cause differences/errors.

### 4. Comparison with other nuclear clusters and origin

To compare the cluster of the Milky Way with other nuclear clusters, we use the size and luminosity measurements of Carollo *et al.* (2002) and Böker *et al.* (2004), which cover mainly early- and late-type spirals, respectively. The size of the cluster is in between these two samples as expected from the Hubble type of the Milky Way. The cluster is however brighter than most other nuclear clusters. That result holds also for the cumulative flux curve which is unaffected by decomposition uncertainties. Our measurement of an usual M/L shows that the cluster of Milky Way has also an unusually high mass. Since its size is typical, this implies that the Milky Way has a higher star density than most nuclear clusters.

The bright nuclear cluster of the Milky Way is like most bright nuclear clusters associated with a bright active circumnuclear region (Launhardt *et al.* 2002, Carollo *et al.* 1999). Seeing that the majority stars in the nuclear cluster are old (Pfuhl *et al.* 2011), this makes it likely also that its old stars have more connections with the nuclear disk than with globular clusters (Tremaine *et al.* 1975). It is however unclear whether a



**Figure 2. Left panel:** Cumulative mass distribution in the GC (with and without the SMBH) in comparison with the literature. The red and orange lines show the results of this work. **Right panel:** The GC in comparison with Carollo *et al.* (2002) and Böker *et al.* (2004). In case of the GC, we show both the size of the nuclear cluster which follows from our two dimensional decomposition (square) and the cumulative flux curve. [A COLOR VERSION IS AVAILABLE ONLINE.]

nuclear disk origin of the nuclear cluster can be reconciled easily with a difference in the flattening in the two components (Section 2).

## References

- Bartko, H., Martins, F., Trippel, S., *et al.* 2010, *ApJ* 708, 834  
 Becklin, E. E. & Neugebauer, G. 1968, *ApJ* 151, 145  
 Beloborodov, A. M., Levin, Y., Eisenhauer, F., *et al.* 2006, *ApJ* 648, 405  
 Blum, R. D., Ramírez, S. V., Sellgren, K., & Olsen, K. 2003, *ApJ* 597, 323  
 Böker, T., Sarzi, M., McLaughlin, D. E., *et al.* 2004, *AJ* 127, 105  
 Carollo, C. M. 1999, *ApJ* 523, 566  
 Carollo, C. M., Stiavelli, M., Seigar, M., de Zeeuw, P. T., & Dejonghe, H. 2002, *AJ* 123, 159  
 Deguchi, S., Imai, H., Fujii, T., *et al.* 2004, *PASJ* 56, 261  
 Dehnen, W. 1993, *MNRAS* 265, 250  
 Do, T., Martinez, G. D., Yelda, S., *et al.* 2013, *arXiv:1311.0886*  
 Genzel, R., Thatte, N., Krabbe, A., Kroker, H., & Tacconi-Garman, L. E. 1996, *ApJ* 472, 153  
 Genzel, R., Eisenhauer, F., & Gillessen, S. 2010, *Reviews of Modern Physics* 82, 3121  
 Gillessen, S., Eisenhauer, F., Trippel, S., *et al.* 2009, *ApJ* 692, 1075  
 Gillessen, S., Eisenhauer, F., Fritz, T. K., *et al.* 2013, *IAU Symposium* 289, 29  
 Haller, J. W., Rieke, M. J., Rieke, G. H., *et al.* 1996, *ApJ* 456, 194  
 Launhardt, R., Zylka, R., & Mezger, P. G. 2002, *A&A* 384, 112  
 Lindqvist, M., Winnberg, A., Habing, H. J., & Matthews, H. E. 1992, *A&AS* 92, 43  
 Matthews, L. D. & Gallagher, J. S., III 1997, *AJ* 114, 1899  
 McGinn, M. T., Sellgren, K., Becklin, E. E., & Hall, D. N. B. 1989, *ApJ* 338, 824  
 Pfuhl, O., Fritz, T. K., Zilka, M., *et al.* 2011, *ApJ* 741, 108  
 Phillips, A. C., Illingworth, G. D., MacKenty, J. W., & Franx, M. 1996, *AJ* 111, 1566  
 Philipp, S., Zylka, R., Mezger, P. G., *et al.* 1999, *A&A* 348, 768  
 Schödel, R., Merritt, D., & Eckart, A. 2009, *A&A* 502, 91  
 Schödel, R. 2011, *The Galactic Center: a Window to the Nuclear Environment of Disk Galaxies* 439, 222  
 Serabyn, E. & Lacy, J. H. 1985, *ApJ* 293, 445  
 Serabyn, E., Guesten, R., Walmsley, J. E., Wink, J. E., & Zylka, R. 1986, *A&A* 169, 85  
 Tremaine, S. D., Ostriker, J. P., & Spitzer, L., Jr. 1975, *ApJ* 196, 407  
 Trippel, S., Gillessen, S., Gerhard, O. E., *et al.* 2008, *A&A* 492, 419  
 Walcher, C. J., van der Marel, R. P., McLaughlin, D., *et al.* 2005, *ApJ* 618, 237